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The dispersion–brightness relation for fast radio bursts from a wide–field survey

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Supplementary Information

1 Data set and data analysis:

Searches were conducted as part of the Commensal Real-time ASKAP Fast Transient³¹ (CRAFT) early science observations (project code AS039). Our search strategy and data analysis largely follow the strategy as reported in the discovery of the first ASKAP-detected FRB⁷.

Searches have been conducted using a number of antennas in the array, with availability limited by commissioning activities and other early-science projects. From March-June 2017, observations were conducted in campaigns using the *main array*: antennas directly interfaced to the telescope correlator and used for commissioning interferometric and synthesis imaging modes for ASKAP. Up to 12 antennas were available for these searches. From June 2017 to present, observations have been conducted using the commissioning subarray, a set of antennas equipped with phased-array feed³² (PAF) and digital subsystems suitable for recording ASKAP-CRAFT fly's eye data³³, but not from the correlator. Over the course of the survey, this sub-array increased in size from 5 to 9 available antennas as additional equipment was integrated.

From 2017 March our search strategy changed from a square beam configuration to a hexagonal close pack configuration (see Figure 1). In both cases the beam separation was 0.9 deg.

Over this time, nearly all of the searches were conducted in 57 fields at Galactic latitudes of $|b| = |50| \pm 5$ deg. Observations of these fields with 55 or 57 min duration were alternated

with 3 or 5-minute scans of known pulsars to check system performance and assess survey efficiency. Fields were assigned automatically by the telescope scheduling software. The selection algorithm was modified on 15 July 2017, to select fields that were close together and at high elevation angle. This was chosen to maximize the sensitivity of a low radiofrequency shadowing campaign with the Murchison Wide Field Array telescope, the results of which will be reported elsewhere. As a consequence of the prioritisation in the selection algorithm, there is large variation in the on-sky time in different fields, with the number of visits for each field ranging from 17 to 1309, with a median of 570.

Searches were conducted with a central frequency of 1320 MHz. The CRAFT data streams are dynamic spectra, produced in the digital beamforming subsystem, providing a total bandwidth of 336 MHz, with 1 MHz spectral resolution and 1.26 ms temporal resolution in pseudo total intensity (i.e., the sum of the XX and YY linearly polarized intensities). These were summed together without any calibration. While beam centre gains between the two polarizations vary by less than⁶ 10%, the beam shapes between the two can vary significantly; calibrating this is an active field of research within the ASKAP commissioning team. Data were recorded using methods nearly identical to those reported previously⁷. Data taking machines were upgraded twice during the course of the searches. Improvements were made to hardware tuning to mitigate network-related failures to record data.

The dynamic spectra were analysed identically to methods presented previously, using the graphical-processor-unit-accelerated search algorithm referred to as *FREDDA*⁷. Improvements were made to the code to better mitigate interference through the subtraction of a nondispersed time series and flag satellite interference. Satellite interference from GPS satellites transmit-

ting in the band is the only significant form of interference persistent throughout the searches. This could be identified as it is confined to a band of frequencies (dependent on to which GPS service the satellite belonged). After flagging for this interference, two impulsive inteferers remained that affected a small fraction (1%) of the data. One was impulsed chirped interference is occasionally produced by diagnostic monitoring of the PAF backplane. It produced candidate bursts with typically high DM ($> 3000 \text{ pc cm}^{-3}$) at the polling interval of 35 s and was eliminated by disabling the monitoring during search observations. Occasional electrical storms would also cause candidates to be detected and would also be identified as episodes of excess low dispersion measure candidates, and could be correlated with on-site weather monitoring systems.

The searches were mostly conducted with on site computing facilities with a typical latency of $< 24 \text{ hr}$ between observation and discovery; a small fraction of the data were searched on a GPU cluster at the Pawsey Centre (<http://pawsey.org.au>), in Perth, with higher latency.

The total exposure is $5.1 \times 10^5 \text{ deg}^2 \text{ h}$. This is the product of the total survey time, effective field of view, and the estimated survey efficiency. The total survey time (counting each antenna individually) was 32,840 hr. To calculate the exposure time we assume an effective field of view of 20 deg^2 , which is the nominal equal power field of view³⁴. This is smaller than the nominal field of view 30 deg^2 , but a more appropriate value for comparison of survey rates because it accounts for the variations in sensitivity across the primary PAF beams. We assume that the observing efficiency is 80%, which is the efficiency of detecting pulses from the pulsars in the 3 – 5 minute hourly scans.

We find no evidence for instrumental or terrestrial origin of the bursts. Bursts were detected in 9 unique antennas, and 19 unique beams, and across a wide range of time-of-day, antenna azimuth and elevation angles, solar angles and seasons. While some 30 deg^2 -fields had multiple detections, they typically had larger-than-average observing times on them (a consequence to the uneven sampling of the fields, as mentioned above). Importantly the bursts show multiple beam detections, and the detections are consistent with arising from the antenna far field. Local interference would be present in more, if not all, of the PAF beams. The brightest burst, FRB180110, shows evidence for an exponential scattering tail, with the tail width varying with frequency as would be expected for propagation through a turbulent plasma. The strongest evidence for the bursts being not produced by interference will be through interferometric localization, a mode which is currently being commissioned.

The candidates are also inconsistent with being produced by thermal noise. A large number of trials have been searched, which we approximate to be $N_{\text{trial}} \approx T_{\text{obs}} R_{\text{samp}} N_b N_{\text{dm}} N_{\text{box}} \approx 1326 \times 7.1 \times 10^6 \times 35 \times 4096 \times 32 \approx 5 \times 10^{17}$, where T_{obs} is the number of antenna days of surveying, R_{samp} is the number of samples per day, $N_b = 35$ is the number of useable beams, $N_{\text{dm}} = 4096$ is the number of trial dispersion measures, and $N_{\text{box}} = 32$ is the number of box-car widths searched. The CRAFT data stream from beam 36 is corrupted and does not produce useful data.

However, the probability of a single 9.5σ candidate arising from thermal noise alone is $P \approx 10^{-21}$. The expected rate of such candidates across the survey is therefore $\approx 5 \times 10^{-4}$. This is likely an overestimate of the probability as the expression for N_{trial} assumes all the trials are independent, which is not the case for DM, width, and time trials.

In addition to strong primary beam detections ($S/N > 9.5$, with the exception of FRB171216, as discussed below). We have chosen to report a S/N of 9.5 as our completeness limit. At this level we were able to verify the FRB nature of the emission. In particular we are able to detect the dispersion of the pulses across the band. Furthermore, we are able to have confidence in the detections as at this level, most of the bursts were detected in multiple beams. 14 of 20 were detected strongly ($S/N > 5$) in other beams, while 19 out of 20 detected at ($S/N > 3$) in adjacent beams.

Below this level, it is still possible to be convinced of the FRB nature of the signals if they have moderate detections in two beams, such as FRB 171216, for which the S/N was greater than 10 when two beams were added together. However we cannot claim that this event is from a complete portion of phase space because our false positive rate increases, so this has not yet been applied to the data set.

The search algorithms have been tested using both simulated data sets and through observations of radio pulsars. We created synthetic filterbank data sets with integration time and bandwidths that were identical to the ASKAP digital system. We generated dispersed pulses over a range of DMs from 0 to 3000 pc cm^{-3} in steps of 25 pc cm^{-3} with an intrinsic width of a single sample. To simulate pulses of differing width, we convolved the pulse in each channel with a Gaussian function of 1 sigma width of 2, 3, and 4 ms, matching the widths of the pulses detected at ASKAP and the majority of those detected at Parkes. We added Gaussian noise at level which yields a S/N of each pulse of 9.5, assuming a matched filter response to the whole pulse. Each dataset of a given DM and width comprised 100 pulses. We searched these data with FREDDA with a S/N threshold of 7. For widths of $\geq 2 \text{ ms}$, we recovered the pulses with a

detection rate of $> 90\%$, and false positive $< 1\%$ for all DMs. For narrower pulses, we expect to lose S/N and recovery due to dilution with the sample integration. Understanding the instrument response to narrow low-DM pulses (narrower than those typically detected at Parkes) would require simulation of voltages of dispersed pulses, channelised by polyphase filters, detected and integrated as in the CRAFT mode. The loss of sensitivity to low-DM narrow pulses is discussed further in supplementary section 5, in relation to differences in Parkes and ASKAP detection rates.

We searched these data with our search algorithm with a S/N threshold of 7 sigma. For widths of ≥ 2 ms, we recovered the pulses with a detection rate of $> 90\%$, and false positive $< 1\%$ for all DMs. For a width of 1 ms (less than a sampling time), the detection rate was 50% at $\text{DM}=75 \text{ pc cm}^{-3}$ and rose to 90% at $\text{DM}=1000 \text{ pc cm}^{-3}$. The loss in sensitivity for pulses narrower than the sampling time is an intrinsic property of any dedispersion search. The optimal delays applied to a given channel should differ depending on where the pulse begins during an integration. The algorithm computes channel delays optimised for a pulse beginning at the center of an integration, thus pulses beginning at the beginning or end of an integration are smeared over two integrations. This effect has the most severe effect for narrow pulses at DMs where DM smearing is not large.

We also have tested the detection pipeline using pulsar data, which was also analysed by a computationally simpler algorithm using a matched filter against the known pulsar properties (pulse width, DM, and approximate Gaussian profile). For the pulses analysed, the pipeline detected the same pulses at $> 90\%$ the S/N reported by the matched filter, at 100% efficiency for all pulses above threshold.

We tested the recovery rate in survey data through hourly observations of the bright pulsars J1644–4559 and J0835–4510. These observations were of 3 or 5 min duration and were processed identically to our high galactic latitude fields. We find that we recover 80% of the pulses, with the remaining 20% lost due to satellite interference. We account for the 20% loss in estimating our total survey time. After automatic flagging of the data for satellite interference, and removing a few observations (totally less than 1% of the survey) affected by the impulsive interference listed above, manual inspection of all the high latitude candidates was conducted, with fewer than 10 candidates produced per day of observation.

Supplementary tables 1 to 19 show the signal-to-noise (S/N) ratios in the primary beam and adjacent beams for FRBs 170416 to 180212. The observations of FRB 170107 were reported previously⁷. We calculate the S/N using the N bins in the frequency-averaged pulse profile, an estimate of σ using off-pulse baseline. In this case

$$S/N = \frac{\sum F_i}{\sigma\sqrt{N}}. \quad (1)$$

Secondary detections and upper limits enabled FRBs to be localized to $\approx 10 \times 10$ arcmin regions, much smaller than the 1.1 deg antenna (primary) beam width for ASKAP at these frequencies. The posterior localization regions are displayed in greyscale relative to the beam positions in Figures 2 to 3, with the intensity of the greyscale proportional to posterior probability; these regions are available in absolute sky coordinates as supplementary files in FITS format (<https://fits.gsfc.nasa.gov/>).

A physical flux scale was applied to data sets by assuming the noise in our data sets is dominated by thermal noise, and the system equivalent flux density S_{sys} is 2000 Jy, consistent

with interferometric measurements of feed sensitivity ⁶ and calibration against simultaneous observations with the Parkes telescope made of PSR B1641–45. In this case the root-mean-square noise is $S_{\text{sys}}/\sqrt{2\Delta\nu\Delta t}$, where $\Delta\nu$ is the spectral bandwidth and Δt is the time resolution. We corrected the integrated pulse profiles (panels labelled *a* in Fig. 1) for off axis attenuation assuming a fixed frequency of 1.3 GHz and the mean offset position of the burst from its primary-beam detection. We did not correct the fluence estimates for a 10% increase in S_{sys} above 1.4 GHz³⁵. This would result in a small correction as most of the band we observe in is below this frequency, and is smaller than the systematic uncertainties by at least a factor of 5. We have however corrected for this effect in the pulse profiles plotted in Figure 1 and in the spectral index estimates.

Burst widths were measured using the frequency integrated spectrum (see Table 1) A non-linear Levenberg-Macquardt fitting algorithm, as implemented in the gnuplot package, was used to measure pulse widths. Most bursts were adequately fitted using a Gaussian function, given the low signal-to-noise ratio of our data. However two bursts (FRBs 180110 and 180130) show evidence for scattering. In this case, the functional form for the function was a Gaussian convolved with an exponential.

Details about specific discoveries are listed below:

FRB 170707: Digital beam weights were not successfully made for some channels, resulting in no signal being recorded for some channels in the burst.

FRB 170906: The data were affected by digitization overflow in the brightest portion of the pulse spectrum. The signal could be recovered by using a beam model and an adjacent beam

detection, where the signal was weaker and therefore not affected by the overflow problem. This was fixed for subsequently detected FRBs. The correction relies on a good understanding of the ASKAP primary beam models distant from the source, so there are larger uncertainties in the peak brightness of the spectrum than for other bursts. One channel of the spectrum in Figure 22 near the peak of bright band drops suddenly; it is possible that this drop is associated with an uncorrected digital artefact.

FRB 171020: The dispersion measure of the burst ($DM=114.1 \pm 0.2 \text{ pc cm}^{-3}$, 1σ) is the lowest detected for an FRB in our sample and lowest for any published burst. The FRB was detected in a corner beam (see supplementary Figure 10); consequently the localization is poor. However, because the extra-Galactic DM is small, the burst likely arises from a nearby galaxy. In addition to the two galaxies listed in the main text, the galaxy cluster Abell 3852 also lies within the FRB uncertainty region. Although there is no published redshift for this cluster, we can estimate its redshift as $z \sim 0.21$ based on the value of $m10 = 19.1 \text{ mag}$ listed by ³⁶ and the empirical relationship between $m10$ and redshift derived by these authors. The galaxies in Abell 3852 are therefore too distant to be plausible hosts for this low-DM FRB.

FRB 171216: The burst was first identified as being an unusually temporally narrow candidate (only one sample in width), but has a $S/N = 8.0$ below the completeness signal-to-noise ratio limit of the survey ($S/N = 9.5$) for a single-beam search. It is therefore not included in tests of the cosmological nature of the distribution. Analysis of adjacent beams show that the burst has nearly equal signal-to-noise ratios in two adjacent beams, as listed in Table 12. The co-added signal to noise ratio from these two detections is 10.3. This co-added pulse is displayed in Figs. 1 and 22.

FRB 180110: The frequency-averaged spectrum of the burst has strong exponential morphology, which would be expected from a pulse scattered off of inhomogeneities in a turbulent plasma. The strength of the scattering (and hence the size of the tail) is highly chromatic. We subdivided the pulse into 8 channels and fitted each independently to a Gaussian convolved pulse of unknown amplitude and width, arrival time, and scattering time. The scattering time changes strongly with frequency (see supplementary Figure 21). We then fitted the scattering times as a function of frequency $\tau(\nu) = A(\nu/\nu_0)^\beta$, and found that $\beta = -3.5 \pm 0.6$ and $A = 4.5 \pm 0.5$ ms at a reference frequency of $\nu_0 = 1.4$ GHz. Our results are consistent with the theoretically predicted values of $\beta = 4.0 - 4.4$ expected for scattering in Kolmogorov media¹⁹ and also observed scattering indices which average to 3.9 ± 0.2 for pulsars in the Milky Way³⁷.

FRB 180119: This burst shows marginal evidence for a scatter broadened tail. The burst is too weak to search for frequency dependent variations as for FRB 180110. The broadening is estimated to be 2.8 ± 0.5 ms.

FRB 180128.2: The data were affected by strong interference in the lower portion of the band. These RFI-affected channels have been flagged in subsequent analyses.

FRB 180130: This burst shows marginal evidence for scatter broadened tail. The burst is too weak to search for frequency dependent variations as for FRB 180110. The broadening is estimated to be 6 ± 1 ms.

2 Analysis of the global spectral properties of the bursts:

We formed spectra for each FRB by integrating the time samples over which the frequency averaged pulse shows detectable signal (nominally at the 1σ noise level). We quote the amplitude of the spectra in fluence, and convert to physical units using an assumed value for the system equivalent flux density ($S_{\text{sys}} = 2000\text{Jy}$, as discussed above) and the measured level of off-pulse noise in the time series using standard methods. The spectra of the bursts can be found in Figure 22 (black lines). The red lines in the same figure show a smoothed version of the spectra. The gray lines show baseline off-pulse spectra taken within 100 ms of each FRB, which provide a guide for the noise level in the FRB spectra and show that data are in general not suffering from excess radio-frequency interference.

Before averaging, we corrected for chromaticity expected due to the beam pattern. This correction is small because the beam pattern densely samples the focal plane, with the beams overlapping well within the half power beam point. We also corrected the spectra for a 10% increase in S_{sys} above 1.4 GHz.

We calculated the average spectra two different ways. In the first, we used natural weighting, in which we directly summed the frequency-dependent fluences for each burst. The average spectra for this can be seen in Figure 23. We fitted a power-law model to the spectra ($E(\nu) \propto \nu^\beta$) using a nonlinear fitting algorithm, and found $\beta = -3.3 \pm 0.3$. We used the bootstrap technique³⁸ to assess how sample variation affects our estimate of the average spectral index. To do this, we formed average spectra, sampling with replacement from the 20 spectra. This formed 1000 composite spectra, each of which was fitted the same power law form for the

flux density. We found that the median spectral index was $\beta = -3.33_{-0.31}^{+0.30}$, matching the fitted value.

In the second method, we weighted the bursts by the inverse of their signal to noise ratio to reduce the bias introduced in the spectra by the bright bursts. When we do this, we find $\beta = -1.8 \pm 0.3$.

In a third method, we formed the spectrum from the median of all 20 FRBs. In this case, we also find $\beta = -1.8 \pm 0.3$.

It is unclear what produces the steep average spectrum. It could be intrinsic to the emission mechanism. Alternatively it could be a secondary propagation effect. In this case, models such as lensing¹⁵ would need to predict a change in spectral occupancy across the band. Diffractive scintillation does not induce steep spectra. Our observations have a relatively small fractional bandwidth of $\approx 10\%$. Broader band observations would strengthen the case for a steep spectral index. Similarly the study of the FRB population at lower frequencies (for example by observing with ASKAP in the lower band or with the other facilities such as the Canadian HI intensity mapping experiment) would also be able to confirm the population has on average a steep spectral index.

3 Source-count analysis:

To analyse the distribution of the bursts, we consider the 19 events above our completeness limit of a signal to noise ratio of 9.5. To examine the source distribution, we use the test statistic^{22,39} V/V_{max} , which is the average fractional volume within which the sources reside

relative to the volume within which they could have been detected. The statistic is calculated from the measured signal-to-noise ratios $((S/N)_i)$:

$$\frac{V}{V_{\max}} = \frac{1}{N_{\text{FRB}}} \sum_i^{N_{\text{FRB}}} \left(\frac{(S/N)_i}{(S/N)_{\min}} \right)^{-3/2}. \quad (2)$$

For a population of sources uniformly distributed throughout Euclidean (flat) space-time, the ensemble average value is $\langle V/V_{\max} \rangle = 0.5$. One advantage of using $V/V_{\max}$ as a test statistic is that it is agnostic to burst width, which varies across the population due to physical and instrumental effects. If there is a correlation between burst width and fluence there could be secondary biases introduced in $V/V_{\max}$ because narrow (but temporally resolved) pulses are easier detect than broader pulses. This will be yet more important when comparing burst populations detected by different telescopes or observing systems. A fuller discussion of the utility of the $V/V_{\max}$ test in the context of FRBs was previously presented²².

The statistic is also independent of the intrinsic luminosity function of the emitting objects, so is valid for both source classes that are standard candles, as well as those that have a broad luminosity function.

For our sample, we find $V/V_{\max} = 0.58$. We calculated the probability that this distribution arises from Euclidean population through simulation. The detected population will depend primarily on the intrinsic burst fluence $E(\nu)$ but secondarily on the width w , and dispersion measure DM. As the purposes of the simulations are to examine a population of local bursts, it is sufficient to assume that these quantities are independent. We assume the population of bursts with fluences that follow a power-law distribution $P(> E(\nu)) \propto E(\nu)^\alpha$. We will assume that the pulses are Gaussian in shape and that the intrinsic width w_i is Normally distributed about some mean μ_w with variance σ_w^2 . We will assume that the dispersion mea-

tures are also Gaussian distributed with a mean value μ_{DM} and variance σ_{DM}^2 . This fit well models the current ASKAP FRB dispersion-measure distribution. We note that our simulations confirm known properties of the V/V_{max} statistic. For example, the standard deviation of the statistic is $\approx 1/\sqrt{12N}$, where N is the size of the population used to test the statistic.

The observed pulse width is instrumentally broadened if the dispersion smearing across a single channel is large or the pulse is narrower than the time sampling t_{samp} . The functional form of this broadening will depend on the impulse response of the spectrometer. We will assume that the effects can be treated as a convolution, and that the impulse response function for both effects are Gaussian as well. In this case the pulse width is

$$w_{\text{obs}} = \sqrt{w_{\text{int}}^2 + w_{\text{DM}}^2 + t_{\text{samp}}^2}, \quad (3)$$

where the dispersion smearing is

$$w_{\text{dm}} = 8.15 \times 10^{-6} \mu\text{s} \left(\frac{\text{DM}}{\text{pc cm}^{-3}} \right) \left(\frac{\Delta\nu}{\text{MHz}} \right) \left(\frac{\nu}{\text{GHz}} \right)^{-3}, \quad (4)$$

where $\Delta\nu$ is the channel bandwidth and ν is the observing frequency.

To model the spectral modulation we multiplied the fluence by a random gain factor g that was exponentially distributed (χ^2 with two degrees of freedom), consistent with diffractive scintillation at high Galactic latitude, where scintillation bandwidths are comparable to the total observing bandwidth⁴⁰.

The observed flux density in time w_{obs} is gE/w_{obs} . Therefore the signal-to-noise ratio is

$$\text{S/N} = \left(\frac{gE}{w_{\text{obs}}} \right) \left(\frac{\sqrt{2Bw_{\text{obs}}}}{S_{\text{sys}}} \right), \quad (5)$$

where B is the total observing bandwidth and the system equivalent flux density S_{sys} factors in aperture efficiency.

To test the consistency of the a Euclidean source-count distribution with the observations, we first generated a large number ($> 10^4$) bursts from that distribution that satisfied our detection S/N. From that sample, we selected (with replacement) 19 events. We then calculated the V/V_{max} statistic for the 19 events. We repeated this process 1000 times, to build up an expected distribution for the test statistic. We then determined the fraction of times the test statistic exceeded the value of $V/V_{\text{max}} = 0.58$ we measured. We note that these simulations assume that the population is both local and homogeneous. For a very nearby population in which sources trace the local large scale structure – closer than a few hundred Mpc so as not to average over the 30 Mpc typical void size – this need not be the case. If the repeating FRB and the non-repeating bursts are in the same population; the distance scale to the former also points to a more distant population. As discussed in the main paper there are lines of evidence suggesting the populations are different. Also as discussed below, the dispersion measure distribution differences between the Parkes and ASKAP suggests a distant population.

Over a wide range of values for the parameters describing the width (μ_w and σ_w) and DM (μ_{DM} and σ_{DM}), we found that for a Euclidean source count distribution, the V/V_{max} values were inconsistent with our measurement at the $\approx 90\%$ confidence level. In particular for values that match the observed DM and (deconvolved) pulse width distributions of the FRBs in our sample ($\mu_{\text{DM}} = 421 \text{ pc cm}^{-3}$, $\sigma_{\text{DM}} = 220 \text{ pc cm}^{-3}$, $\mu_w = 2.5 \text{ ms}$, $\sigma_w = 1.6 \text{ ms}$), we find the probability that the fluence distribution is inconsistent with Euclidean is 88%, using a one-sided test (i.e., $P(V/V_{\text{max}} > 0.58)$). A one sided test is appropriate as values of $V/V_{\text{max}} < 0.5$

are not expected. The distribution of simulated test statistic (black line) assuming a Euclidean energy distribution (black line) and its expected distribution based on the assumption that $\sigma = 1/\sqrt{12N}$ are displayed in supplementary Figure 24. In supplementary Figure 25, we show the results of simulations of pulses with ASKAP-like DM distributions derived from fluence distributions with $\alpha = -1.5$ (Euclidean) and $\alpha = -2.1$. The ASKAP fluences produce fewer bright FRBs than expected from a Euclidean distribution, and better match the steeper source counts. The top panels (*a* and *d*) show the DM distributions of the ASKAP (blue) and Parkes-detected FRBs (red), in addition to the simulated ASKAP population (blue). As can be seen the simulated and observed ASKAP distributions match well.

To test which values of source count slope(α) were consistent with the observed value of the test statistic, we simulated event lists at a range of source count distributions and re-computed the $V/V_{\max}$ distribution. We found the value of α for which $V/V_{\max} > 0.58$ 50%, 17%, and 83% of the time, and used these values to define the maximum likelihood and 67% confidence region for α .

The value of the $V/V_{\max}$ test statistic does not change significantly when setting a higher threshold. When setting the threshold to be $S/N > 15$, we find $V/V_{\max} = 0.66 \pm 0.10$, consistent with that calculated at a S/N of 9.5.

4 FRB fluence dispersion-measure distribution:

In Figure 2, we have plotted the FRBs detected at Parkes^{1,2,5,16,17,25,28,41–47} at a frequency of 1.4 GHz; those detected at the upgraded Molonglo Telescope (UTMOST^{21,48–51}) at a central

frequency of 0.843 GHz; and at the Green Bank telescope²⁰ at a frequency of 0.8 GHz. For the repeating FRB (121102), we have plotted both the detected fluence for FRB 121102⁵² at a frequency of 1.4 GHz, as well as the fluences of repeated bursts detected at observation frequencies close to 1.4 GHz^{3,29}. We have assumed published uncertainties are accurate for fluence lower limits. As there is large uncertainty in burst location within the beam pattern for non-ASKAP FRB discoveries, the upper limits are quoted at twice the reported fluence. We note that observations of the repeating FRB after it was discovered have more reliable measurements and in Figure 2, we show the reported error bars. For Parkes FRBs 150807 and 110724 we have highlighted fluences corrected by modelling multiple beam detections. FRB 110724 has the high but uncertain fluence $E_\nu = 800 \pm 400$ Jy ms; uncertainties are dominated by instrumental effects.

The burst dispersion measures have been corrected for Galactic contributions using the NE2001 model⁵³ for the electron density of the Milky Way Disk. We also remove a nominal Milky Way halo contribution, a practice consistent with the DM calculation of the repeating FRB⁴. We empirically estimate a likely halo contribution of approximately 15 pc cm^{-3} , based on the lowest dispersion measures of pulsars in the Large Magellanic Cloud pulsar survey⁵⁴; the minimum DMs are 65, 68 and 69 pc cm^{-3} and NE2001 reports a Galactic contribution in this direction of 52 pc cm^{-3} . For most halo models (e.g., based on X-ray absorption measurements⁵⁵) the decline in the density beyond the $\sim 50 \text{ kpc}$ distance of the LMC provides at most a further $1\text{-}2 \text{ pc cm}^{-3}$ beyond this (the DM contribution at larger radii, r , is substantially smaller than the further contribution to the halo mass since each shell adds mass proportional to r^2). Other estimates of the halo contribution⁵⁶ are within the $\sim 10 \text{ pc cm}^{-3}$ uncertainty of the Milky Way DM contribution.

The conversion between dispersion measure and redshift has been calculated assuming concordance cosmology and a smooth distribution of electron density¹⁰, and assuming that the host dispersion measure contribution is $50(1+z)^{-1} \text{ pc cm}^{-3}$. Bursts are likely to show large variations in host-galaxy DM contribution, with variations depending on the viewing geometry. At the mean viewing angle, host galaxy contribution to the DM has been estimated to be $30 - 40 \text{ pc cm}^{-3}$ for elliptical and dwarf galaxies and $50 - 60 \text{ pc cm}^{-3}$ for spiral galaxies⁹.

5 Comparing the DM distributions from Parkes and ASKAP

There is a notable difference in the DM distributions from Parkes and ASKAP. This difference is intrinsic to the FRB populations and not due to instrumental selection effects.

This can be demonstrated through the same simulation approach implemented for the source-count analysis. The modern Parkes system, which now dominates the Parkes detections has higher time ($\Delta t_{\text{samp}} = 64 \mu\text{s}$) and frequency resolution ($\Delta\nu = 390 \text{ kHz}$) than the ASKAP system ($\Delta t_{\text{samp}} = 1.26 \text{ ms}$, $\Delta\nu = 1 \text{ MHz}$). As a result, high dispersion measure pulses have larger instrumental broadening at ASKAP.

We use the same simulation procedures we used to examine the statistics of the V/V_{max} distribution to test the sensitivity of the system to pulses of differing widths and DM. We considered pulses with Euclidean source-count distribution. Instead of considering events with pulse widths and DM consistent with an ASKAP-like distribution, we use broader distributions to examine how selection bias could be affecting the ASKAP distributions. Therefore, we simulated pulses with uniform probability density of arising between 50 and 4000 pc cm^{-3} , in

$\log(\text{DM})$. The purpose of this DM distribution is not to simulate a physical DM distribution but to allow plots of the detected DM distribution to be easily interpreted. Histograms of the distribution of the detected pulses in $\log(\text{DM})$ directly show how sensitivity is lost with DM. Pulses were assumed to have widths between 0.1 and 10.0 ms with equal probability density in $\log(\text{width})$.

While there is a difference in total rate detected by the two systems, there is no evidence for ASKAP losing a disproportionate number of high-DM events. Figure 26 shows the normalised dispersion measure distributions for events detected by ASKAP-like (blue) and Parkes-like (red) backend systems. The figure shows that above a DM of $\approx 300 \text{ pc cm}^{-3}$ the shapes are identical. This indicates that above this DM, both systems are losing events in equal proportion. Above the Parkes diagonal DM, the largest diagonal DM, both systems sensitivity decrease at the same rate with respect to DM. Therefore above this DM, any difference in the DM distribution reflects intrinsic properties of the bursts detected by the two telescopes. The only difference between the two systems is that the Parkes-like system would be expected to detect more low-DM events (a consequence of the higher time resolution of Parkes), which is also in complete disagreement with the observed DM distributions of the two telescopes.

6 Interpreting the non-repetition of the ASKAP sample

To limit the repetition properties of ASKAP FRBs, we assume that they all repeat with identical rates R :

$$R(S > S_0) = R_0(S/S_0)^\gamma. \quad (6)$$

Here, R_0 is the rate above the ASKAP detection threshold S_0 (nominally 46 Jy assuming a width close to the median of 4 ms). We note that the detection depends on pulse width. This scales to other fluence thresholds S according to the cumulative repetition rate index γ .

The probability of detecting no repeat pulses from any of the observed FRBs can be calculated as:

$$p(\text{no repeats}) = \prod_{i=1}^{N_{\text{FRB}}} (1 + R_0 T_i B_i^\gamma) e^{-R_0 T_i B_i^\gamma}, \quad (7)$$

where the terms in the product represent the Poisson probability of observing 0 or 1 bursts given a T_{obs} of T_i , and beam correction factor B_i (i.e. $S_0 \propto B^{-1}$). This allows ASKAP FRB repetition rates greater than once per 62 (35) [23] days to be excluded with 68% (95%) [99.7%] confidence in a one-sided test for $\gamma = -1$. This value was chosen because it provides the steepest pulse-energy distribution consistent with the repeating FRBs, and thus provides the most conservative constraint on repetition within the ASKAP sample.

Equation 6 is used to scale to the sensitivity at which FRB 121102 has been observed, comparing to previous VLA observations²⁶, for which a 7.4σ threshold to each spectral window with 1σ sensitivity of 5 mJy over 5 ms produces $S_{\text{VLA}} = 184 \text{ mJy ms}$. The resulting limits on $R(184 \text{ mJy ms})$ are shown in Figure 27.

Law et al.²⁶ report $\gamma_{\text{VLA}} = -0.6_{-0.2}^{+0.3}$. Rates are taken from the *late-2016* and *early-2016* samples, reporting 9 (nominal uncertainty ± 3 , i.e. $0.32 \pm 0.1 \text{ hr}^{-1}$) and 0 bursts, respectively. We exclude almost the entire *late-2016* region in R - α space at 99.7% confidence, with a small region at $\gamma = -0.9$ being allowed. Performing a joint maximisation of R to allow for fluctuations in both the observed VLA and ASKAP rates results in $R(184 \text{ mJy ms}) = 0.18 \text{ hr}^{-1}$ for

$\gamma = -0.9$, for which the probability product of ASKAP seeing no repeat bursts, and the VLA seeing at least nine, is 0.4%. This demonstrates that ASKAP FRBs do not exhibit repetition with the same characteristics as FRB 121102.

Comparisons between FRB 121102 and the limits on the repetition properties of ASKAP FRBs presented here should be taken with caution, due to the different observation bands, frequency structure, pulse duration, and dispersion measures. In particular, equation 7 would be more correctly formulated in terms of intrinsic source properties, so that non-observations of repeats from nearby (presumably low-DM) FRBs would be more significant. A more complete analysis would also include a model of the undetected repeating FRBs, since non-observations of multiple pulses from these will further limit the repetition rate through equation 7. Finally, note that any deviations from Poisson statistics in pulse arrival times will lead to clustering, and hence stronger limits on repetition for those FRBs which are necessarily detected in their active phase.

We note that our constraints on the repetition rate use a relatively high threshold for other bursts, comparable to the one used for blind FRB searches. Searches at lower thresholds are possible because they can focus on a narrow dispersion measure range in the pointings directed toward the FRBs. Sub-threshold searches are ongoing with effort being placed to characterize false alarms due to interference at these lower thresholds.

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Supplementary Tables

Supplementary Table 1: FRB 170416 detections

Beam	RA (J2000)	DEC (J2000)	S/N
	(deg)	(deg)	
15	332.24	-54.834	7.8
16	333.03	-55.62	8.3
20	333.01	-54.06	2.0
21	333.80	-54.83	12
27	334.54	-54.04	2.0
28	335.36	-54.81	2.0

Supplementary Table 2: FRB 170428 detections

Beam	RA (J2000)	DEC (J2000)	S/N
	(deg)	(deg)	
19	326.92	-41.84	8.2
20	327.53	-42.62	0.9
14	326.31	-42.62	2.5
13	325.71	-41.83	1.4
18	326.33	-41.06	2.0
25	327.52	-41.06	1.3
26	328.13	-41.84	0.6

Supplementary Table 3: FRB 170707 detections

Beam	RA (J2000)	DEC (J2000)	S/N
	(deg)	(deg)	
00	44.17	−56.58	5.2
01	44.89	−57.39	10.5
02	45.64	−58.19	3.4
06	45.80	−56.64	3.3
07	46.56	−57.44	2.7

Supplementary Table 4: FRB 170712 detections

Beam	RA (J2000)	DEC (J2000)	S/N
	(deg)	(deg)	
16	339.02	−60.67	13.8
09	337.18	−60.66	0.5
10	338.09	−61.45	6.4
15	338.11	−59.89	2.4
17	339.97	−61.44	6.9
21	339.91	−59.88	0.7
22	340.86	−60.65	1.9

Supplementary Table 5: FRB 170906 detections

Beam	RA (J2000)	DEC (J2000)	S/N
	(deg)	(deg)	
20	329.77	−19.80	18.6
19	330.39	−20.49	10.1
14	330.71	−19.64	8.6
21	329.16	−19.11	3.5
26	329.44	−20.64	7.4
27	328.83	−19.95	3.7

Supplementary Table 6: FRB 171003 detections

Beam	RA (J2000)	DEC (J2000)	S/N
	(deg)	(deg)	
06	187.96	−13.05	0.1
18	186.74	−14.06	8.1
12	187.65	−13.90	14.3
13	187.05	−13.22	3.4
19	186.14	−13.38	0.0
25	185.82	−14.22	0.0
24	186.42	−14.91	1.1

Supplementary Table 7: FRB 171004 detections

Beam	RA (J2000)	DEC (J2000)	S/N
	(deg)	(deg)	
26	179.29	−12.07	11.4
25	179.89	−12.76	1.6
31	178.98	−12.91	1.4
32	178.39	−12.22	2.3
27	178.71	−11.38	3.1
20	179.61	−11.22	4.8
19	180.20	−11.91	4.7

Supplementary Table 8: FRB 171019 detections

Beam	RA (J2000)	DEC (J2000)	S/N
	(deg)	(deg)	
21	334.59	−8.34	24.3
20	335.17	−9.03	9.1
15	335.48	−8.18	0.5
16	334.90	−7.49	1.1
22	334.00	−7.65	4.4
28	333.69	−8.49	12.1
27	334.27	−9.18	17.9

Supplementary Table 9: FRB 171020 detections

Beam	RA (J2000)	DEC (J2000)	S/N
	(deg)	(deg)	
00	333.22	−19.99	24.4
01	332.59	−19.31	4.0
06	332.27	−20.16	0.0
07	331.65	−19.48	1.8
12	331.96	−21.01	0.9

Supplementary Table 10: FRB 171116 detections

Beam	RA (J2000)	DEC (J2000)	S/N
	(deg)	(deg)	
28	53.34	−17.35	11.9
22	53.66	−16.51	1.8
21	54.27	−17.20	0.0
27	53.94	−18.04	0.3
33	53.00	−18.19	4.6
34	52.40	−17.50	6.0
29	52.74	−16.65	8.1

Supplementary Table 11: FRB 171213 detections

Beam	RA (J2000)	DEC (J2000)	S/N
	(deg)	(deg)	
01	54.99	−11.97	2.4
02	54.39	−11.28	27.1
03	53.80	−10.60	6.8
07	54.08	−12.13	2.0
08	53.49	−11.44	2.0

Supplementary Table 12: FRB 171216 detections

Beam	RA (J2000)	DEC (J2000)	S/N
	(deg)	(deg)	
20	49.91	−57.47	0.6
21	50.77	−58.25	0.3
26	50.72	−56.69	1.5
27	51.58	−57.46	8.4
28	52.48	−58.22	-0.6
32	52.36	−56.66	8.6
33	53.25	−57.42	3.4

Supplementary Table 13: FRB 180110 detections

Beam	RA (J2000)	DEC (J2000)	S/N
	(deg)	(deg)	
30	327.44	−34.30	−0.1
31	328.00	−35.08	34.7
32	328.57	−35.85	30.6
33	329.15	−36.62	0.1
25	326.90	−35.08	2.2
beam26	327.46	−35.86	16.7

Supplementary Table 14: FRB 180119 detections

Beam	RA (J2000)	DEC (J2000)	S/N
	(deg)	(deg)	
19	52.86	−13.14	9.1
20	52.27	−12.45	17.3
21	51.68	−11.76	2.0
14	53.17	−12.29	3.3
15	52.58	−11.60	2.3
26	51.95	−13.29	9.0
27	51.36	−12.60	2.7

Supplementary Table 15: FRB 180128 detections

Beam	RA (J2000)	DEC (J2000)	S/N
	(deg)	(deg)	
11	209.01	−6.73	13.0
05	209.90	−6.58	3.4
10	209.59	−7.42	4.2
17	208.70	−7.58	1.7

Supplementary Table 16: FRB 180128.2 detections

Beam	RA (J2000)	DEC (J2000)	S/N
	(deg)	(deg)	
03	335.35	−60.63	10.8
02	334.53	−59.83	6.6
04	336.21	−61.43	−2.1
08	336.32	−59.87	7.3
09	337.18	−60.66	4.2

Supplementary Table 17: FRB 180130 detections

Beam	RA (J2000)	DEC (J2000)	S/N
	(deg)	(deg)	
22	328.05	−38.19	10.8
21	327.47	−37.42	-0.5
23	328.65	−38.97	6.7
16	326.91	−38.20	2.3
17	327.49	−38.98	7.0
28	328.61	−37.41	1.7
29	329.20	−38.18	1.8

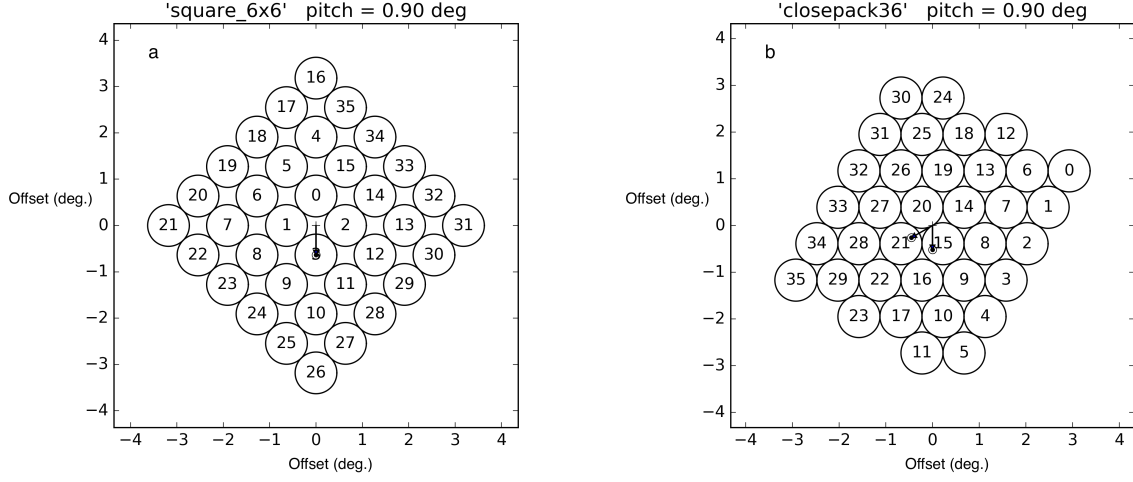
Supplementary Table 18: FRB 180131 detections

Beam	RA (J2000)	DEC (J2000)	S/N
	(deg)	(deg)	
25	327.52	−41.06	14.8
24	326.92	−40.28	9.3
30	328.10	−40.28	9.0
31	328.71	−41.05	1.9
26	328.13	−41.84	−1.0
19	326.92	−41.84	0.2
18	326.32	−41.06	3.5

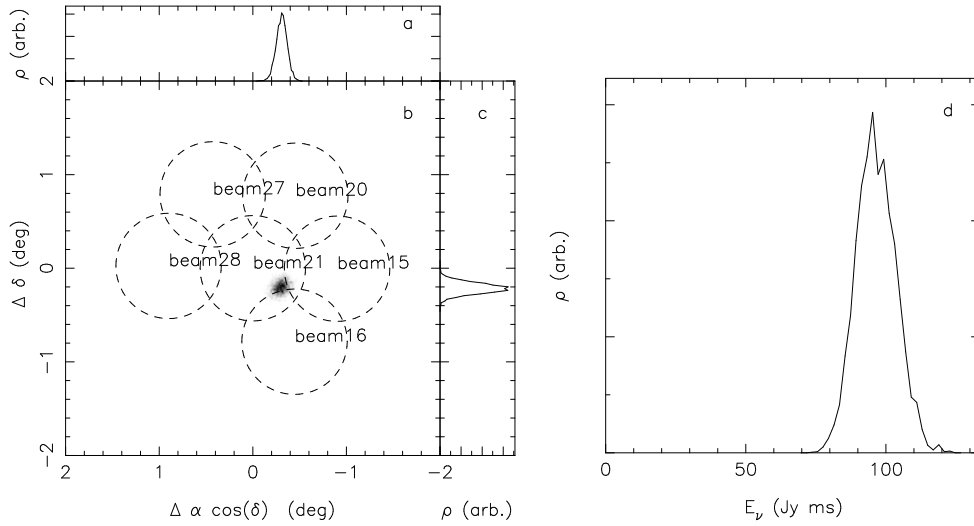
Supplementary Table 19: FRB 180212 detections

Beam	RA (J2000)	DEC (J2000)	S/N
	(deg)	(deg)	
05	215.27	−3.80	18.3
04	215.85	−4.49	4.7
10	214.96	−4.64	0.9
11	214.38	−3.95	3.1

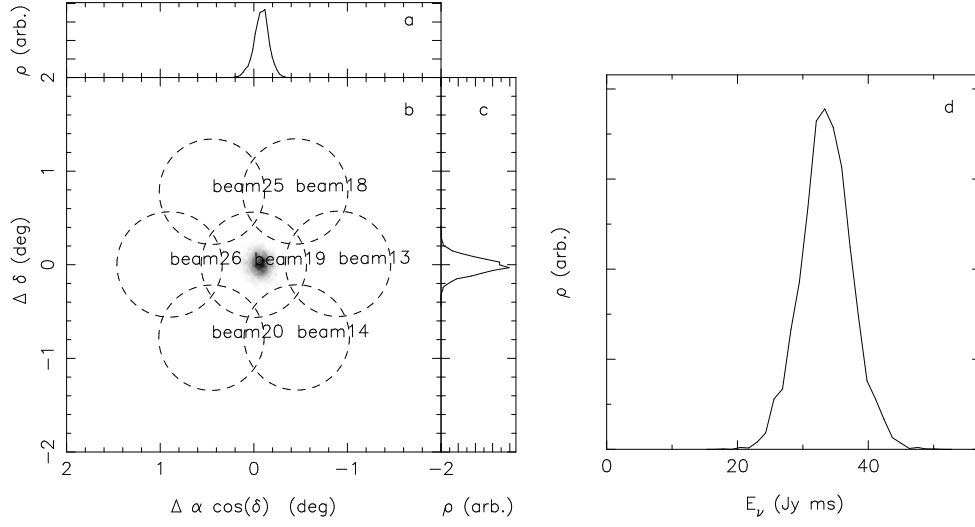
Supplementary Figures



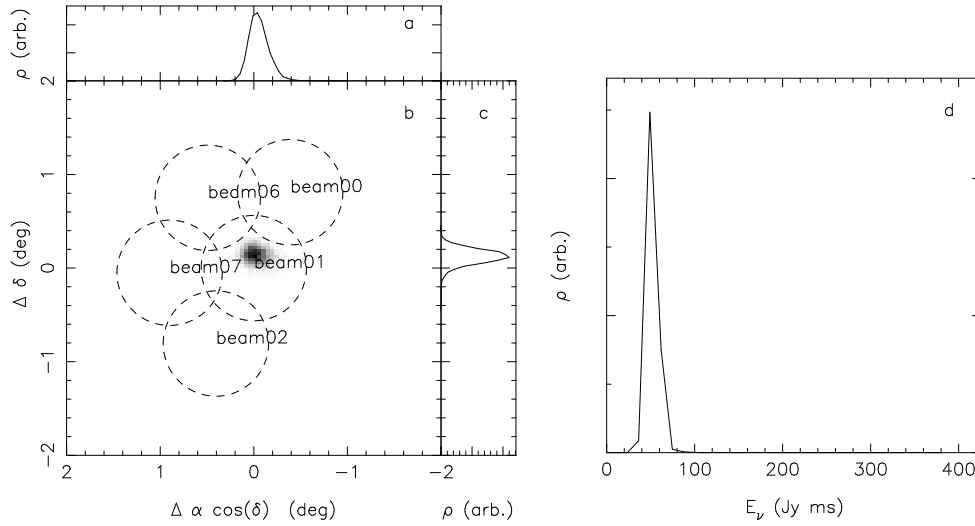
Supplementary Figure 1: Phase array beam configurations. In panel a, we show the square configuration used for initial survey that discovered FRB 170107. In panel b, we show the hexagonal configuration used in subsequent searches. The circle diameters are 0.9 degree, slightly smaller than the FWHP beam point at these frequencies⁶, which is $\approx 1.2 \text{ deg}(\nu/1.3 \text{ GHz})^{-1}$.



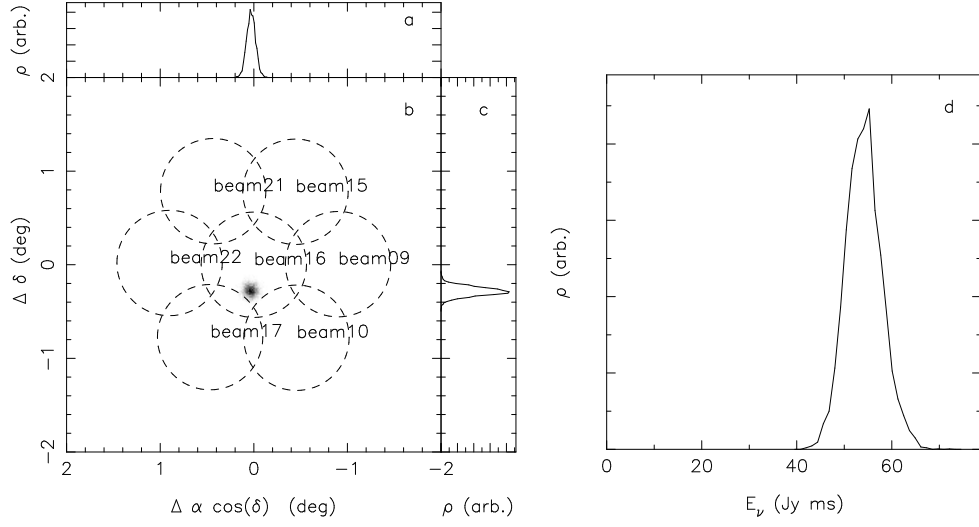
Supplementary Figure 2: Constraints on position and fluence for FRB 170416. Panels a, b and c show the one and two-dimensional posterior distributions for burst position. Panel d shows the constraint on burst fluence.



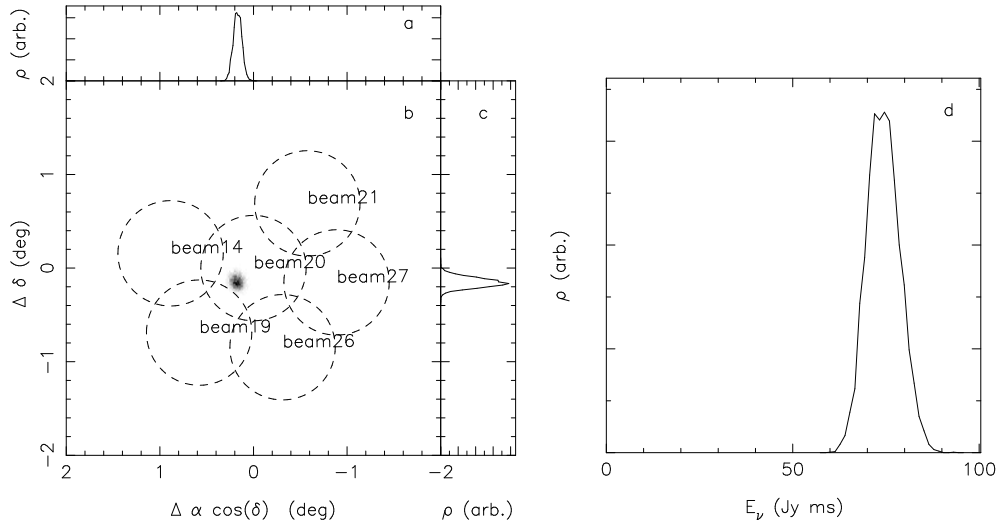
Supplementary Figure 3: Constraints on position and fluence for FRB 170428. Panels a, b and c show the one and two-dimensional posterior distributions for burst position. Panel d shows the constraint on burst fluence.



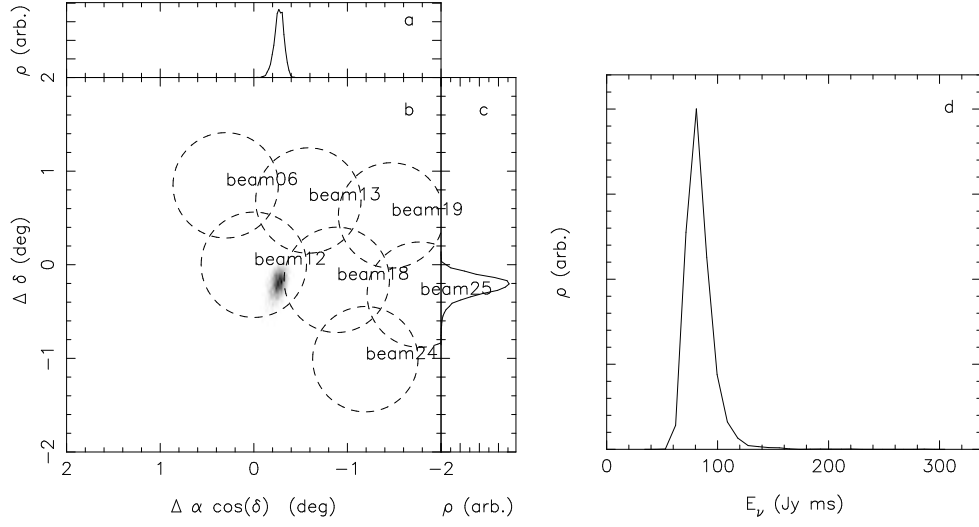
Supplementary Figure 4: Constraints on position and fluence for FRB 170707. Panels a, b and c show the one and two-dimensional posterior distributions for burst position. Panel d shows the constraint on burst fluence.



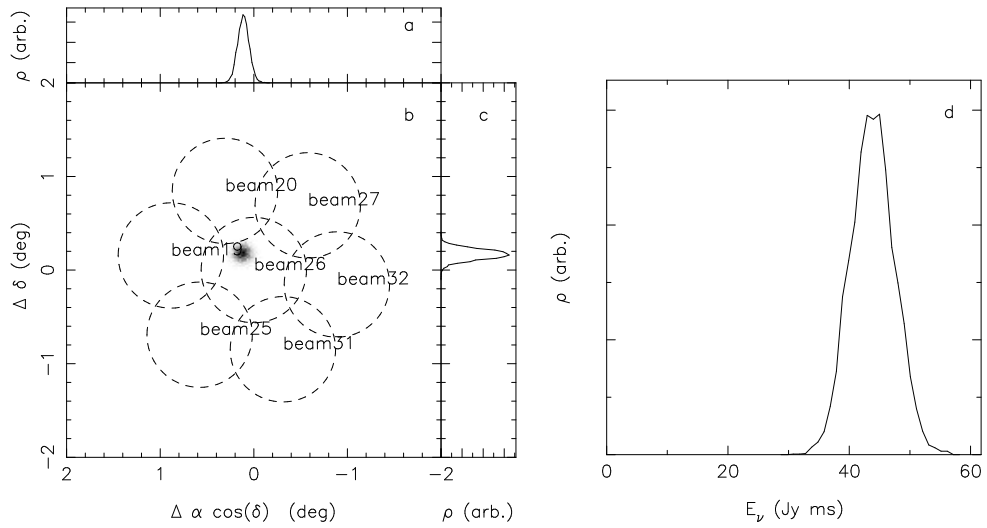
Supplementary Figure 5: Constraints on position and fluence for FRB 170712. Panels a, b and c show the one and two-dimensional posterior distributions for burst position. Panel d shows the constraint on burst fluence.



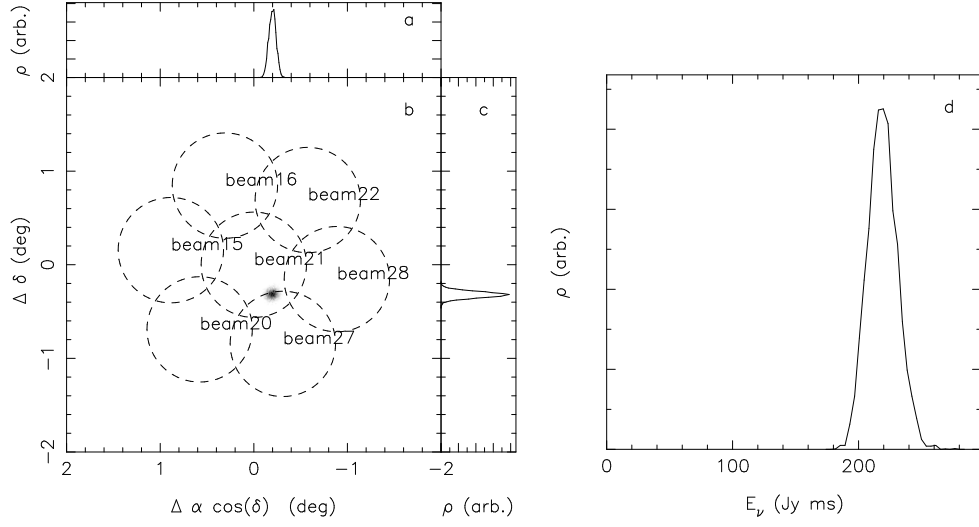
Supplementary Figure 6: Constraints on position and fluence for FRB 170906. Panels a, b and c show the one and two-dimensional posterior distributions for burst position. Panel d shows the constraint on burst fluence.



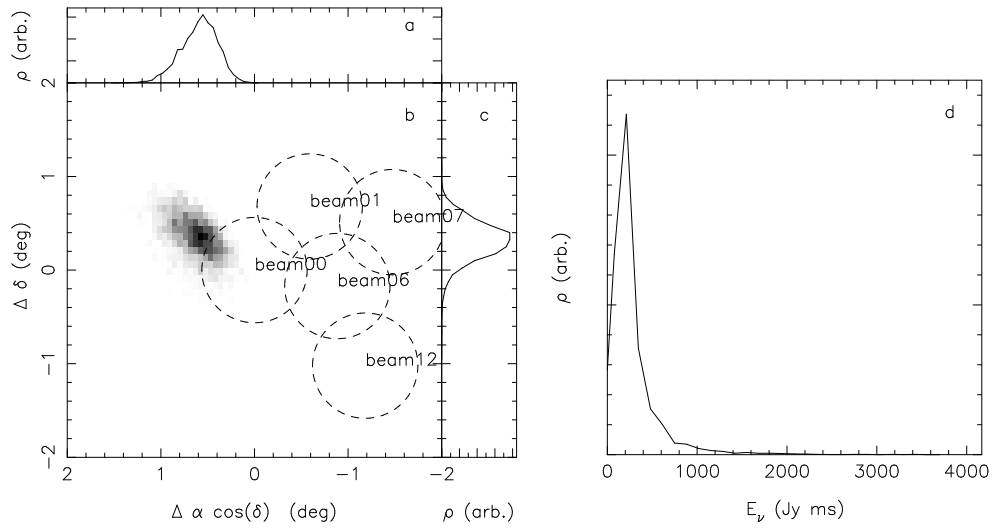
Supplementary Figure 7: Constraints on position and fluence for FRB 171003. Panels a, b and c show the one and two-dimensional posterior distributions for burst position. Panel d shows the constraint on burst fluence.



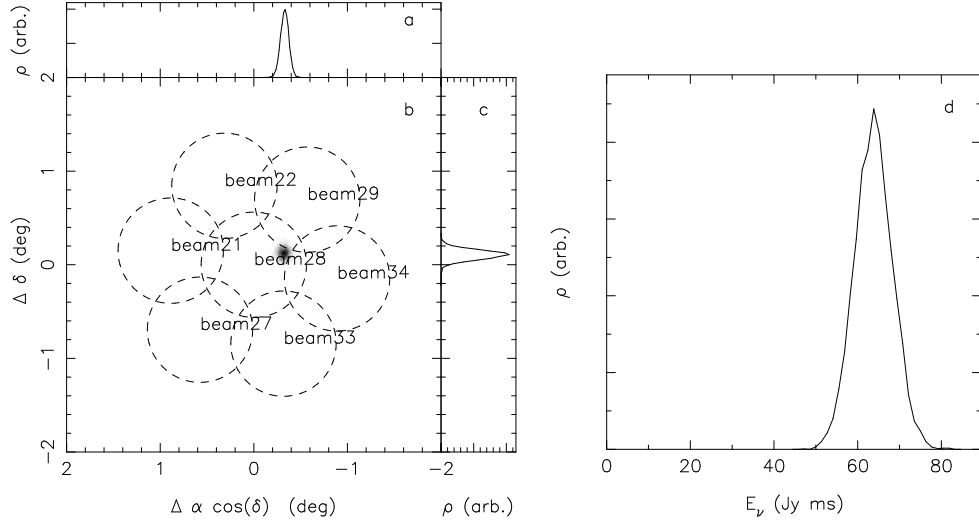
Supplementary Figure 8: Constraints on position and fluence for FRB 171004. Panels a, b and c show the one and two-dimensional posterior distributions for burst position. Panel d shows the constraint on burst fluence.



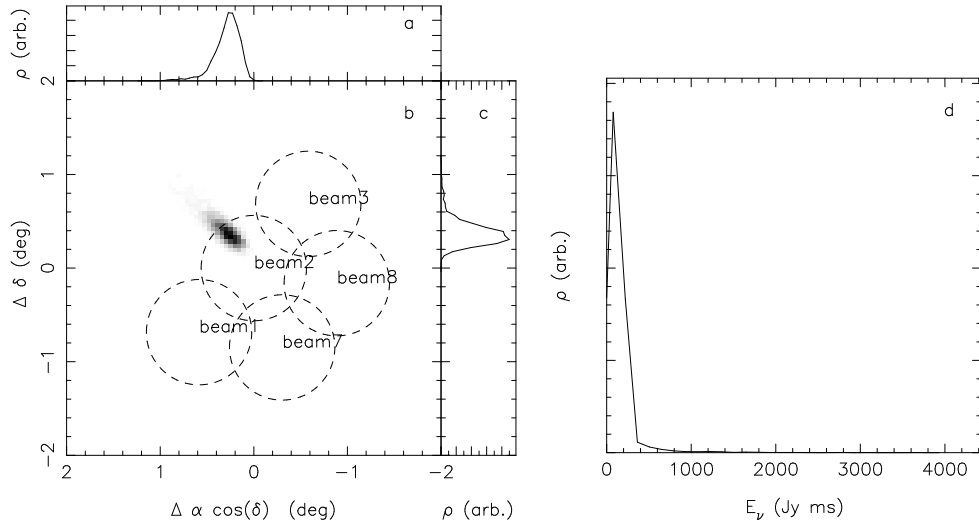
Supplementary Figure 9: Constraints on position and fluence for FRB 171019. Panels a, b and c show the one and two-dimensional posterior distributions for burst position. Panel d shows the constraint on burst fluence.



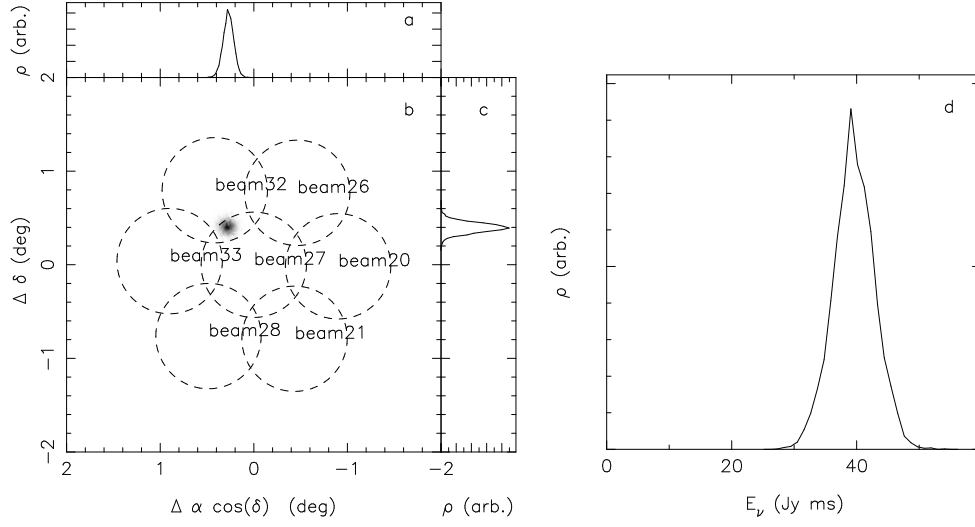
Supplementary Figure 10: Constraints on position and fluence for FRB 171020. Panels a, b and c show the one and two-dimensional posterior distributions for burst position. Panel d shows the constraint on burst fluence.



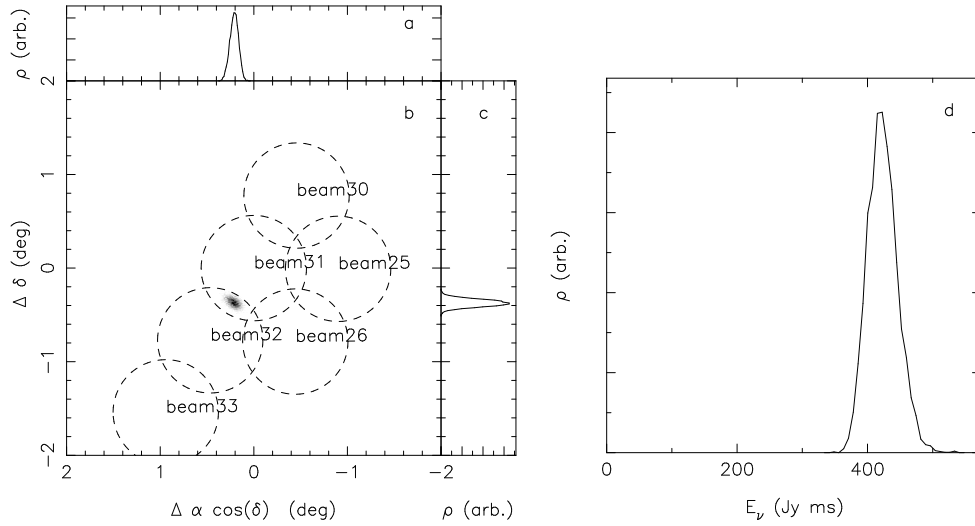
Supplementary Figure 11: Constraints on position and fluence for FRB 171116. Panels a, b and c show the one and two-dimensional posterior distributions for burst position. Panel d shows the constraint on burst fluence.



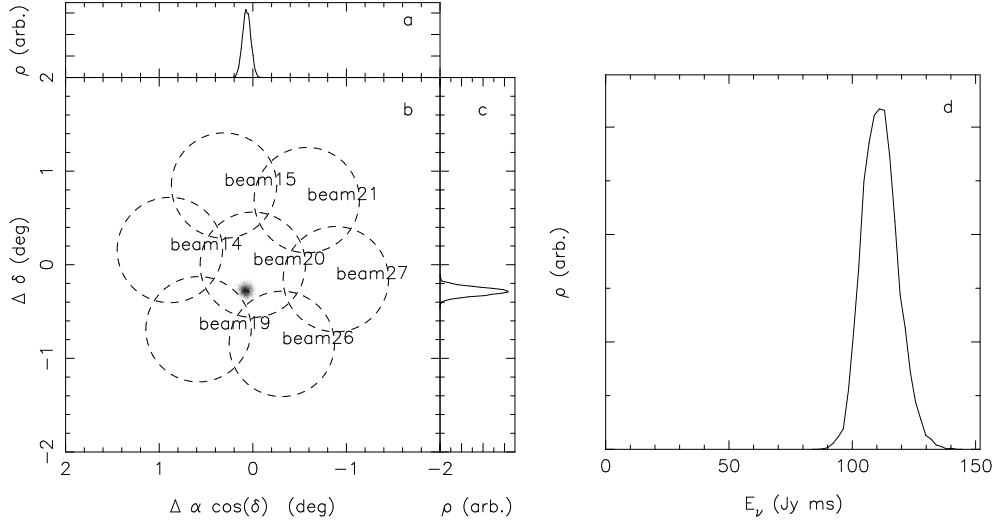
Supplementary Figure 12: Constraints on position and fluence for FRB 171213. Panels a, b and c show the one and two-dimensional posterior distributions for burst position. Panel d shows the constraint on burst fluence



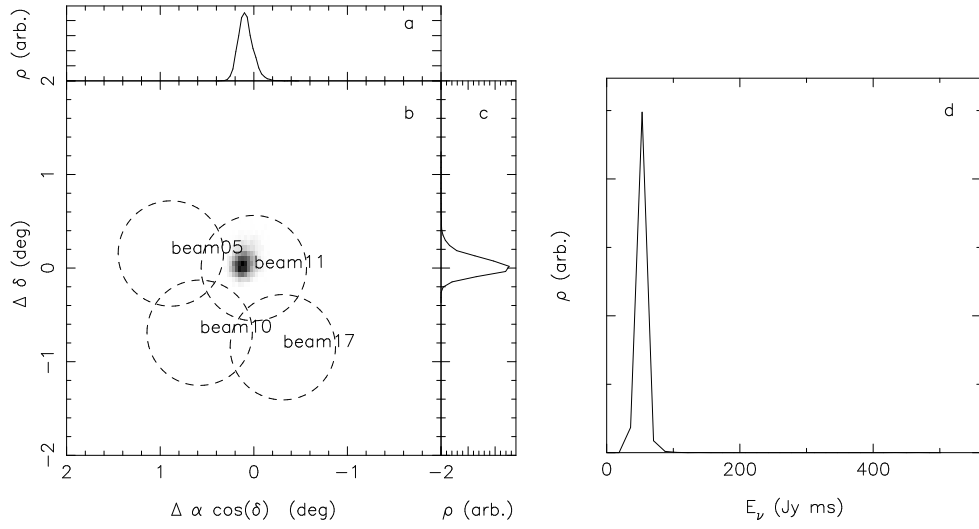
Supplementary Figure 13: Constraints on position and fluence for FRB 171216. Panels a, b and c show the one and two-dimensional posterior distributions for burst position. Panel d shows the constraint on burst fluence



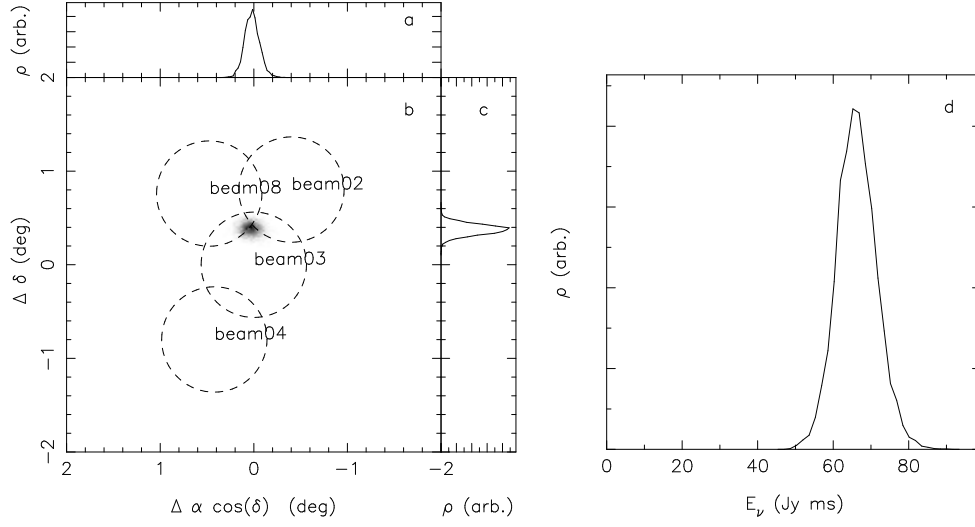
Supplementary Figure 14: Constraints on position and fluence for FRB 180110. Panels a, b and c show the one and two-dimensional posterior distributions for burst position. Panel d shows the constraint on burst fluence



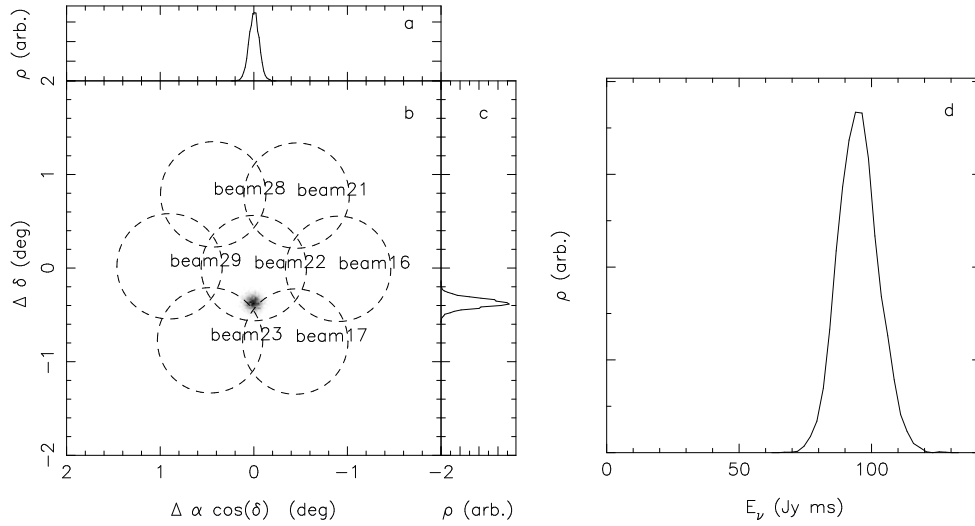
Supplementary Figure 15: Constraints on position and fluence for FRB 180119. Panels a, b and c show the one and two-dimensional posterior distributions for burst position. Panel d shows the constraint on burst fluence



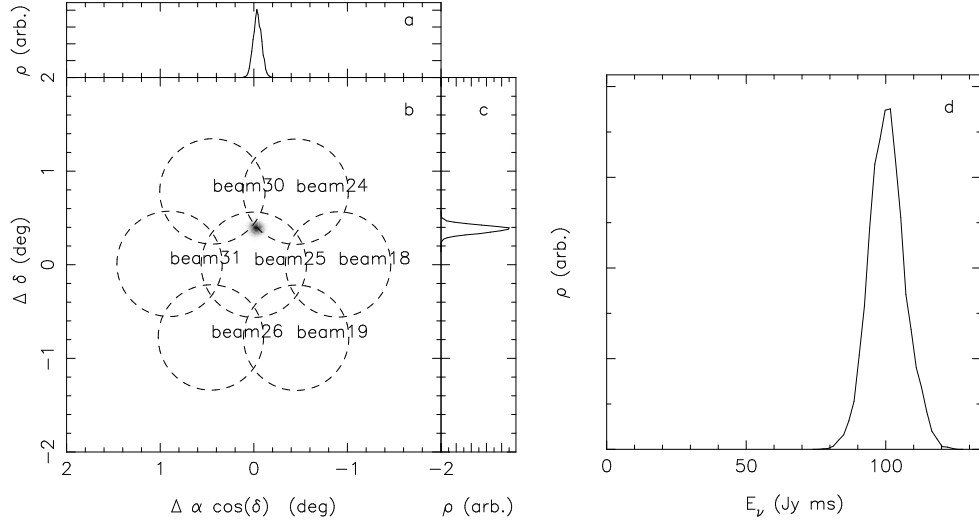
Supplementary Figure 16: Constraints on position and fluence for FRB 180128.0. Panels a, b and c show the one and two-dimensional posterior distributions for burst position. Panel d shows the constraint on burst fluence



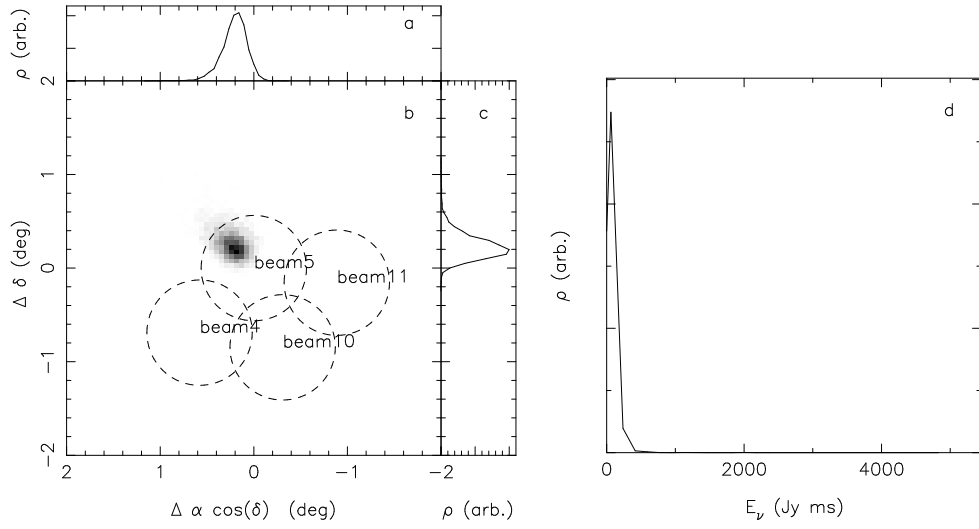
Supplementary Figure 17: Constraints on position and fluence for FRB 180128.2. Panels a, b and c show the one and two-dimensional posterior distributions for burst position. Panel d shows the constraint on burst fluence



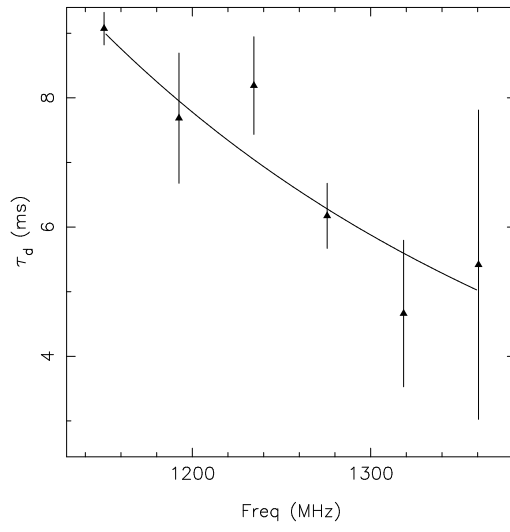
Supplementary Figure 18: Constraints on position and fluence for FRB 180130. Panels a, b and c show the one and two-dimensional posterior distributions for burst position. Panel d shows the constraint on burst fluence



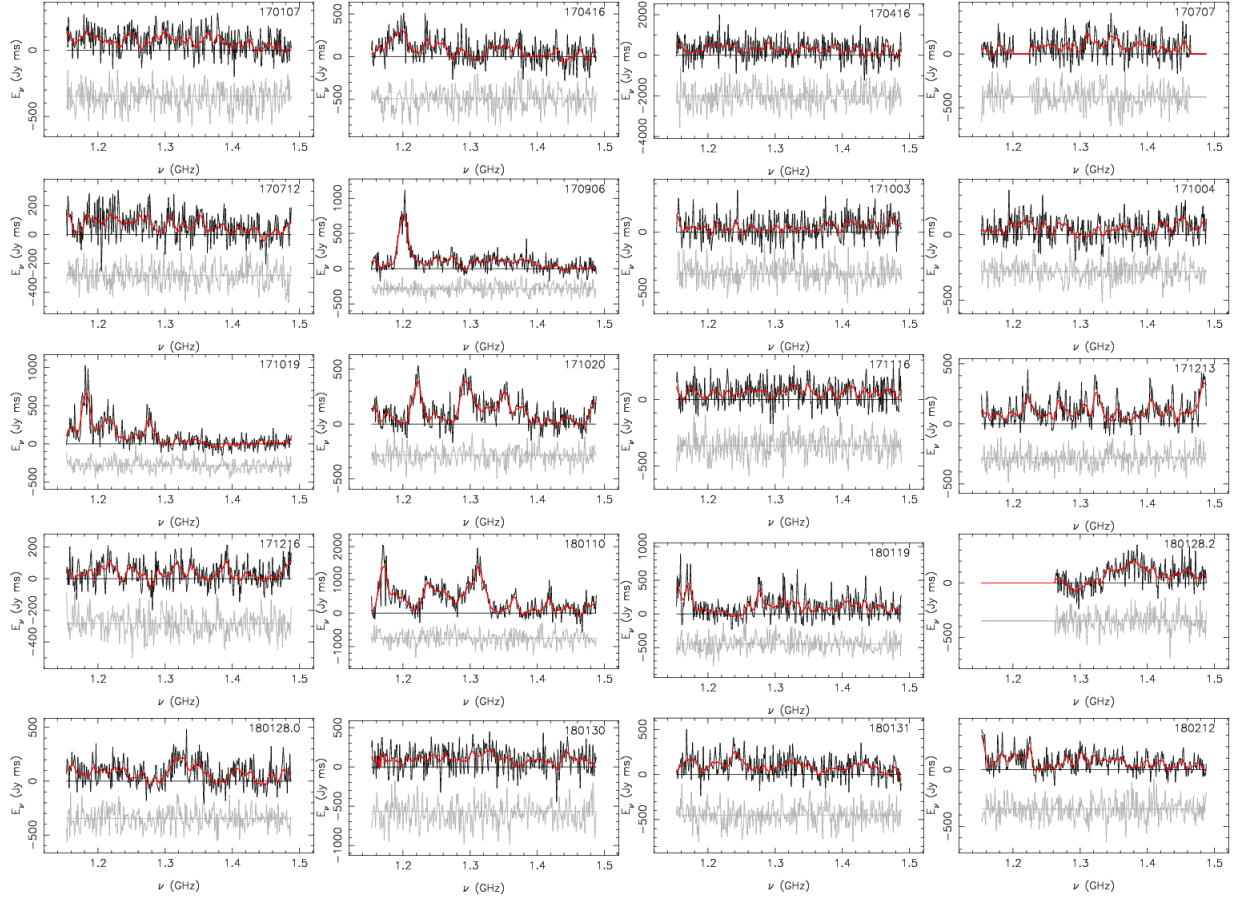
Supplementary Figure 19: Constraints on position and fluence for FRB 180131. Panels a, b and c show the one and two-dimensional posterior distributions for burst position. Panel d shows the constraint on burst fluence



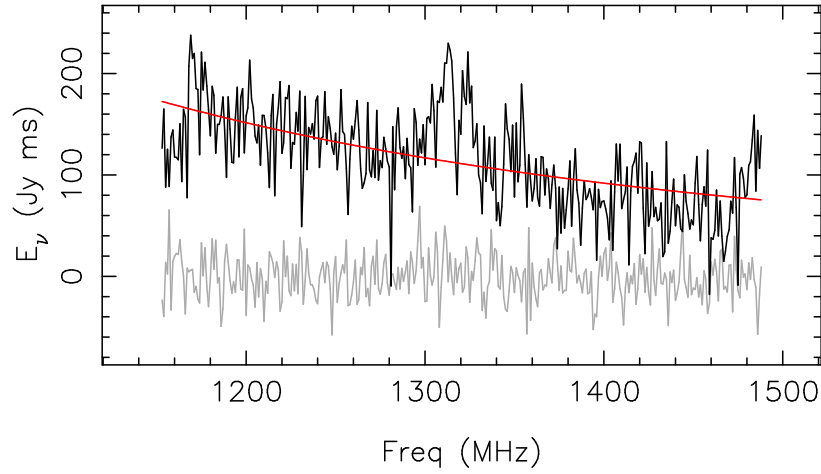
Supplementary Figure 20: Constraints on position and fluence for FRB 180212. Panels a, b and c show the one and two-dimensional posterior distributions for burst position. Panel d shows the constraint on burst fluence



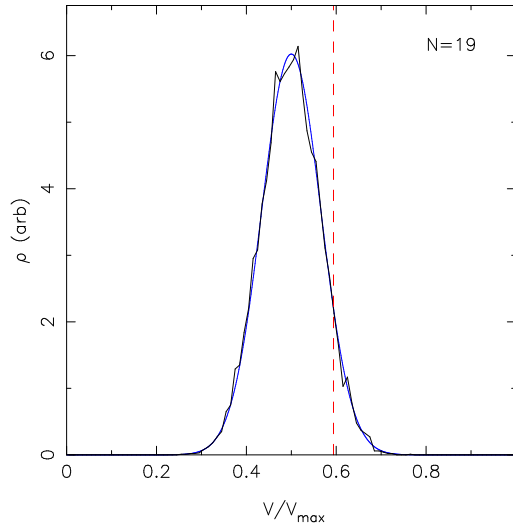
Supplementary Figure 21: Pulse broadening for FRB 180110. The broadening time τ_d shows significant change with frequency ν . The solid line shows the best-fitting power law as a function of frequency $\tau_d(\nu) = (4.5 \pm 0.5 \text{ ms}) (\nu/\nu_0)^{-3.5 \pm 0.5}$, where the quoted uncertainties are 1σ .



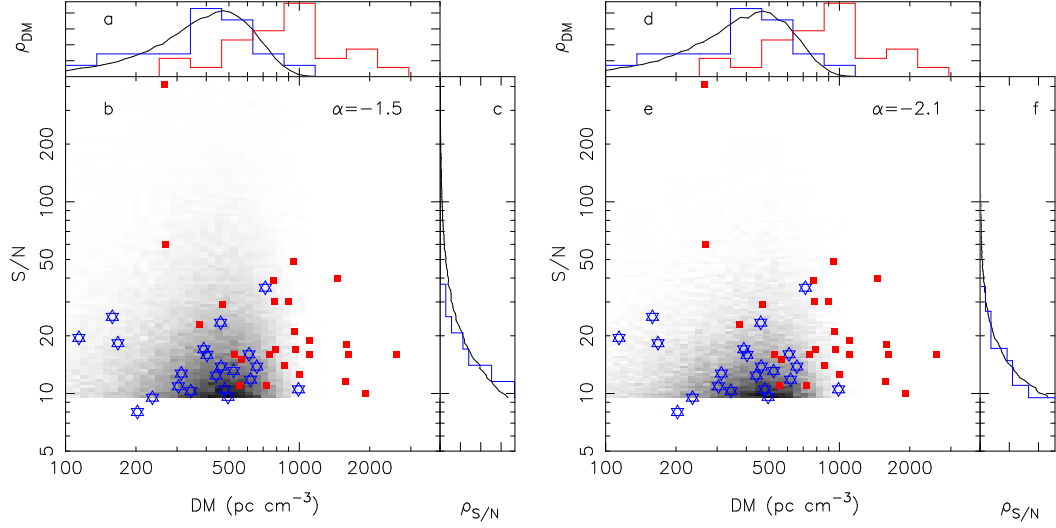
Supplementary Figure 22: Spectra for ASKAP FRBs. Spectra are produced from the primary beam detection, with the exception of FRB 171216, and have not been corrected for primary beam gain variations. For FRB 171216 the spectra has been produced by adding together the two 8σ detections with equal weighting. Red lines are smoothed spectra, smoothed using a Gaussian Kernel with standard deviation of 4 MHz (4 channels). Grey lines are example baseline spectra taken within 100 ms of the FRBs, and are offset from zero for clarity. Horizontal lines show zero power for both the on and off pulse spectra. For FRB 170707, a portion of the spectrum is missing due to spectral channels with zero beam weights (see supplementary text). For FRB 180128.2, the zeroed channels are flagged due to radio frequency interference (see supplementary text).



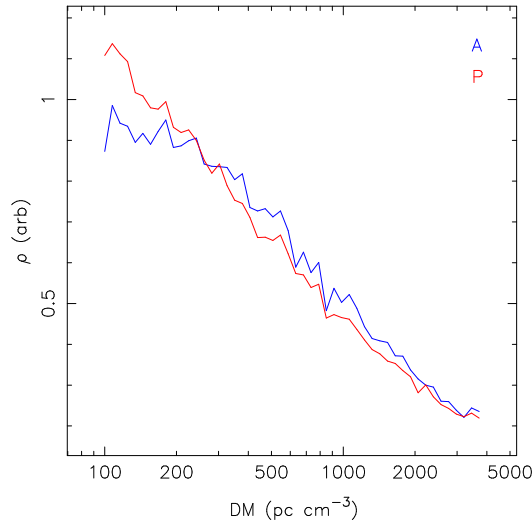
Supplementary Figure 23: Average spectrum of the ASKAP-detected FRBs. The spectrum is shown in black; in red, we show the best fitting power law (red). The gray spectrum shows the summed and dedispersed off-pulse spectra taken from within 100 ms of the FRB events. The average power-law spectra index is -1.8 ± 0.3 .



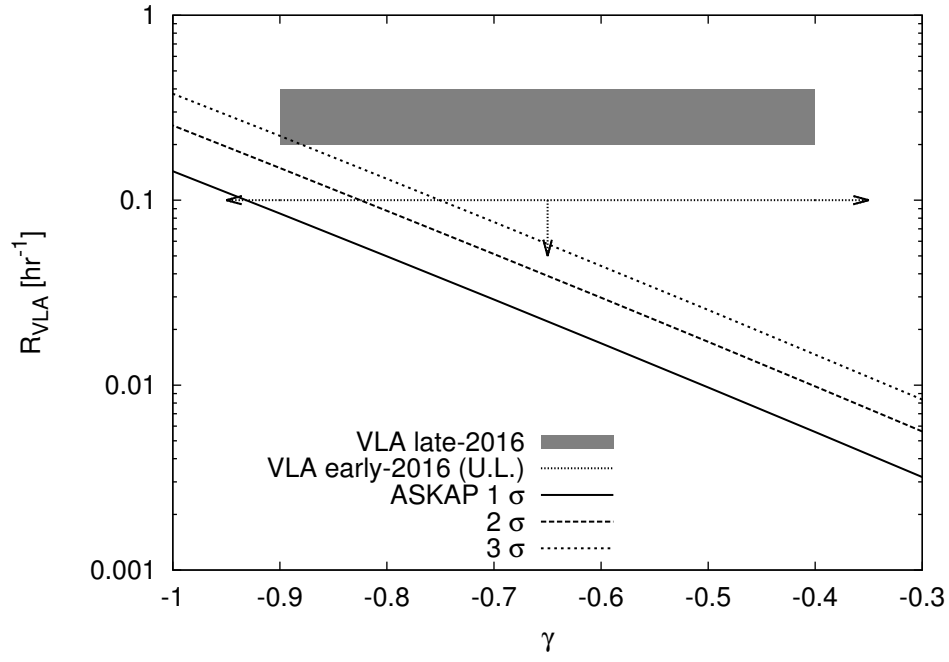
Supplementary Figure 24: Distribution of the $V/V_{\max}$ test statistic. The black line shows the simulated distribution. The blue curve is an exponential distribution with standard deviation $\sigma = 1/\sqrt{12N}$, with $N = 19$. The red dashed line is the value of test statistic derived from the 19 bursts detected above our completeness limit.



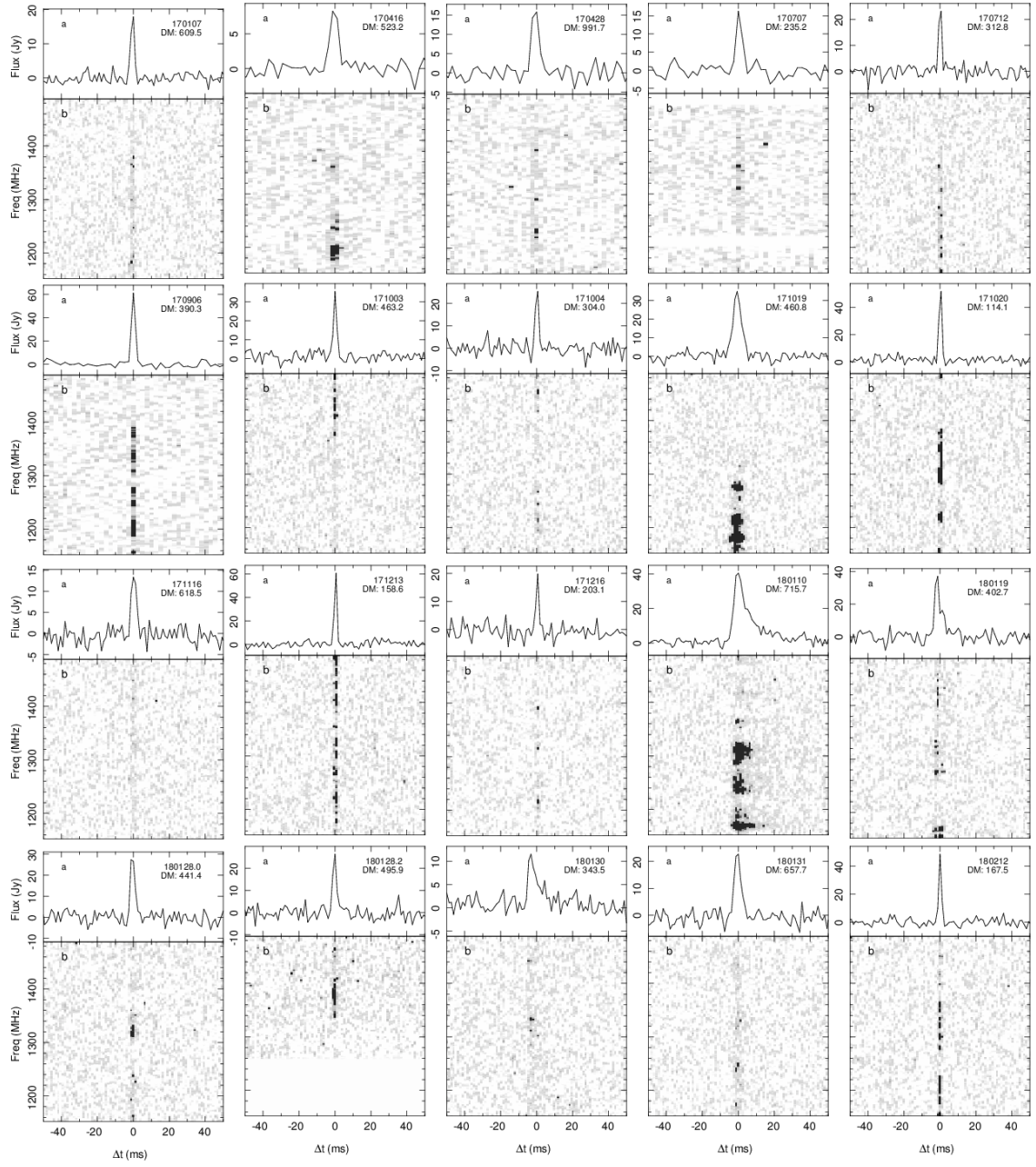
Supplementary Figure 25: Fluence-DM distributions for FRBs. In panels *a-c* we show the simulated distributions for $\alpha = -1.5$. In panels *d-f* we show distributions for $\alpha = -2.1$ (right). Panels *a* and *d*: Dispersion-measure distribution. The black curve shows the simulated DM distribution. The blue histogram is the ASKAP DM distribution. The red histogram is the Parkes DM distribution. The simulations match the ASKAP DM distribution. Panels *b* and *e*: Joint distribution of S/N and DM. The grayscale shows the 2-D simulated distribution. The ASKAP FRBs are shown as blue stars. The Parkes FRBs are shown as red squares. Panels *c* and *f*: S/N distribution. The black curve shows the S/N distribution from simulation. The blue curve shows the S/N distribution from the ASKAP events. A Euclidean distribution produces an excess of bright events relative to what we observe.



Supplementary Figure 26: Simulated response of receiving systems to dispersed pulses. The plot shows the response of ASKAP-like (blue) and Parkes-like (red) backend receiving systems to a power-law population of FRBs with a Euclidean fluence distribution. If the systems were ideal, (infinite time and frequency resolution) the distributions would be flat. At high DM ($> 800 \text{ pc cm}^{-3}$) there is no difference in the shape of the distributions, indicating that the observed difference in median DM is not the result of the poorer time and frequency resolution of the ASKAP data-recording instrument.



Supplementary Figure 27: Comparison of repetition of FRB 121102 to limits from ASKAP. The plot shows the limits on the repetition rate R of ASKAP-detected FRBs at 184 mJy ms sensitivity as a function of the cumulative repetition rate index α , compared to observations of FRB 121102 by the JVL²⁶.



Supplementary Figure 28: Pulse profiles and dynamic spectra of ASKAP fast radio bursts. The FRB names and dispersion measures (DM) in units of pc cm^{-3} are listed in the panels showing the pulse profiles (Panels a). The spectra (Panels b) have been dedispersed to the maximum-likelihood dispersion measure. The colour scale is set to range from the mean to 4σ of the off-pulse intensity. Time averaged spectra can be found in Supplementary Figure 22.